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A TEST OF PRE-MAIN-SEQUENCE LITHIUM DEPLETION MODELS

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ABSTRACT

Despite the extensive study of lithium depletion during pre-main-sequence (PMS) contraction, studies of individual stars show discrepancies between ages determined from the Hertzsprung–Russell (H–R) diagram and ages determined from lithium depletion, indicating open questions in the PMS evolutionary models. To further test these models, we present high-resolution spectra for members of the β Pictoris Moving Group (BPMG), which is young and nearby. We measure equivalent widths of the 6707.8 Å Li I line in these stars and use them to determine lithium abundances. We combine the lithium abundance with the predictions of PMS evolutionary models in order to calculate a lithium depletion age for each star. We compare this age to the age predicted by the H–R diagram of the same model. We find that the evolutionary models underpredict the amount of lithium depletion for the BPMG given its nominal H–R diagram age of ~12 Myr, particularly for the mid-M stars, which have no observable Li I line. This results in systematically older ages calculated from lithium depletion isochrones than from the H–R diagram. We suggest that this discrepancy may be related to the discrepancy between measured M-dwarf radii and the smaller radii predicted by evolutionary models.

Key words: open clusters and associations: individual (β Pictoris Moving Group) – stars: abundances – stars: activity – stars: evolution – stars: late-type – stars: pre-main sequence

Online-only material: color figures

1. INTRODUCTION

Lithium depletion during pre-main-sequence (PMS) contraction has been extensively studied. Studies of open clusters have shown that lithium depletion is a strong function of both age and stellar mass (e.g., Soderblom et al. 1993; Barrado y Navascués et al. 2004; Mentuch et al. 2008). These studies clearly show that the presence of a strong Li I line at 6707.8 Å is an indicator of youth in late-type stars. The regularity of lithium depletion in open clusters implies that the lithium line can act as a mass-dependent clock (Jeffries & Naylor 2001). By combining a measurement of the lithium abundance of a star with theoretical evolutionary models, one might be able to derive an age of a PMS star whose distance is unknown, e.g., a young field star. In order to do this, we need to understand PMS lithium depletion very well.

Most of the work comparing lithium depletion to Hertzsprung–Russell (H–R) diagram ages has been done in open clusters, focusing on comparing the cluster age derived from the lithium depletion boundary (LDB)³ in the lowest-mass stars with the H–R diagram age found from fitting the upper main-sequence (e.g., Barrado y Navascués et al. 1999; Stauffer et al. 1999; Burke et al. 2004; Jeffries & Oliveira 2005; Manzi et al. 2008). Most studies find that the LDB age tends to be older than the age calculated from upper main-sequence fitting, though the discrepancy may be smaller for younger clusters (Jeffries & Oliveira 2005; Manzi et al. 2008). These results are used to argue for the inclusion of convective overshooting in evolutionary models, since it increases the ages derived from upper main-sequence fitting without affecting the LDB age. These older upper main-sequence fitting ages tend to be in better agreement with LDB ages (Burke et al. 2004).

³ For stars that are fully convective before they reach the main sequence (M < 0.3 M☉), at a given age there is a sharp boundary between stars that have fully depleted their lithium and slightly less massive stars with no evidence of depletion (Basri et al. 1996; Bildsten et al. 1997).
begin by presenting our spectra and lithium equivalent widths (EWs) in Section 2. We then determine effective temperatures, luminosities, and lithium abundances for the stars in our sample in Section 3. In Section 4, we compare these data with the H–R diagram and lithium depletion isochrones of theoretical models. We find that the models systematically underpredict the observed lithium depletion for M stars (or equivalently, overpredict their lithium depletion ages compared to H–R diagram ages). We suggest that the observed tendency for models of main-sequence M stars to underpredict stellar radii may hold the key to resolving the discrepancy between the observed and predicted lithium depletion in late-type stars. We conclude in Section 6.

2. OBSERVATIONS AND DATA

We took spectra of ten of the late-K and M stars or binary systems listed in Table 1 of Zuckerman et al. (2001), with separate spectra for each of the components of the binary stars GJ 799 and BD $-17^\circ6128$; the spectra are shown in Figures 1 and 2. We also took a spectrum of a somewhat earlier-type member of the BPMG, the spectroscopic binary (SB) HD 155555 A/B. The spectra were taken with the 4 m Blanco telescope at Cerro Tololo Inter-American Observatory on six consecutive nights between 2001 October 30 and 2001 November 5, using the echelle spectrograph. The spectrograph covered a wavelength range from 4800 Å to 8400 Å with a measured spectral resolving power of $R \sim 40,000$. The spectra were reduced in IRAF\(^4\) according to standard procedures. We flux calibrated the spectra with respect to two flux standards, HR 9087 and HR 1544, using standard, sensfunc, and calibrate in IRAF. We visually examined the flux calibrated spectra to

\(^4\) IRAF is distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.
ensure that the calibration was consistent between orders and between the two calibration standards.

2.1. Spectral Type

To determine spectral types for our sample, we compared the strength of the TiO5 band to the strength of the nearby continuum, following the method given in Reid et al. (1995). We calculated the ratio of the flux from 7126 Å to 7135 Å to the flux from 7042 Å to 7046 Å. Using the linear fit relation from Reid et al. (1995), we took the ratio of these fluxes to determine the spectral types of our sample stars. According to Reid et al. (1995), this method gives an uncertainty of ±0.5 spectral subtype. Note that this relation is not well defined for spectral types much earlier than M0 because the TiO band heads are very weak or disappear entirely. Thus, the spectral types measured for HIP 29964, BD −17°6128 A, and CD −64°1208 are more uncertain, although probably correct in a relative sense. Additionally, since the relation gives a numeric value for the spectral type, with 0 representing M0, it is not clear how to assign that number to the Morgan–Keenan system for values ≤−1 since some authors argue that K6 is not a distinct spectral type (e.g., Kirkpatrick et al. 1991). We follow the convention of omitting K6 as a spectral type and assign −1 to K7 and −2 to K5. Thus, we give the spectral types of HIP 29964 and BD −17°6128 A as K5.4 and K5.7, respectively, corresponding to −1.6 and −1.3 as measured from the relation given in Reid et al. (1995). Zuckerman et al. (2001) give the spectral types of these stars as K6/7 and K7/M0, respectively, and Torres et al. (2006) find K4 and K6. The spectral types of the stars later than K7 agree within the errors with the types given in Zuckerman & Song (2004), with the exception of HD 155555 C, for which Zuckerman & Song (2004) give M4.5. However, our value of M3.5 for this star is in agreement with the M3 spectral type found by Torres et al. (2006). BD −17°6128 B was not listed independently in Zuckerman & Song (2004), but an estimate of its spectral type is given in Neuhäuser et al. (2002) as M0–2, which is close to the spectral type of M2.9 that we find there.

2.2. Equivalent Width of Li i

Using the splot routine in IRAF, we measured the EWs of all identifiable Li i lines at 6707.8 Å (EW(Li i)).

3. ANALYSIS

3.1. Luminosity and Rotational Velocity

We determined the luminosity and projected rotational velocity (v sin i) of each star based on data from the literature. We calculated the luminosity of each star using V magnitudes from Table 1 in Zuckerman et al. (2001). We assigned a fixed value of σV = ±0.1 as the error on the V magnitudes to account for measurement errors and variability. We used bolometric corrections from Kenyon & Hartmann (1995) appropriate to the spectral type of the star. Distances to each star came from the new reduction of the Hipparcos catalog (van Leeuwen 2007). The V magnitude of BD −17°6128 is the magnitude of the combined system, so we adopt the luminosities given in Neuhäuser et al. (2002) for the A and B components.

For the v sin i of each star, we used the average of the values given in Zuckerman et al. (2001), Torres et al. (2006), and Scholz et al. (2007). For CD −64°1208, we used the value given in Torres et al. (2006), v sin i = 110 km s−1, as Zuckerman et al. (2001) simply give v sin i ≥ 100 km s−1. BD −17°6128 B does not have a v sin i given in the literature. We measured the v sin i by artificially broadening a template star spectrum of similar spectral type (Figure 1). We find that the v sin i of BD −17°6128 B is 20 ± 5 km s−1.

3.2. Effective Temperature

We converted the measured spectral type for each star into an effective temperature T eff using empirical spectral type—T eff relations to calculate the effective temperature of each star. For spectral types M1–M9, we used the temperature scale given in Luhman (1999) for “intermediate” stars, since our young stars can be expected to fall in this class. We combined this with the table given in Kenyon & Hartmann (1995) for earlier spectral types and interpolated to get effective temperatures for our stars. Changes within the spectral type uncertainties lead to typical changes in effective temperature of ± ~100 K. These data are given in Table 1.

3.3. Lithium Abundances

We calculated lithium abundances (A(Li) = 12 + log(N(Li)/N(H)); Jeffries 2000) for our stars using the curves equivalent width. For simplicity, we use the term “equivalent width” to refer to both equivalent widths (as measured for K and earlier spectral types) and pseudo-EW (as measured for the M dwarfs).

Figure 2. Spectrum of CD −64°1208 showing the Li i (6707.8 Å) and Ca i (6717.7 Å) lines. Note that each line appears to have two components, shown by tick marks at measured RVs of −28 and 43 km s−1.
Table 1  Late-type Members of the BPMG

<table>
<thead>
<tr>
<th>Name</th>
<th>Other Name</th>
<th>Spectral Type</th>
<th>Distance (pc)</th>
<th>$V$ (mag)</th>
<th>$v \sin i$ (km s$^{-1}$)</th>
<th>$T_{\text{eff}}$ (K)</th>
<th>$L_\star$ ($L_\odot$)</th>
<th>EW(Li$i$) (mÅ)</th>
<th>A(Li)</th>
<th>Notes</th>
</tr>
</thead>
<tbody>
<tr>
<td>HD 155555 A</td>
<td>V824 Ara</td>
<td>G5 IV</td>
<td>31.4 ± 0.5</td>
<td>7.21</td>
<td>34</td>
<td>5770$^{+50}_{-40}$</td>
<td>1.1 ± 0.1</td>
<td>120 ± 20</td>
<td>...</td>
<td>1, 2, 4</td>
</tr>
<tr>
<td>HD 155555 B</td>
<td>V824 Ara</td>
<td>K0 IV/V</td>
<td>31.4 ± 0.5</td>
<td>8.08</td>
<td>33</td>
<td>5250$^{+125}_{-140}$</td>
<td>0.54 ± 0.06</td>
<td>250 ± 60</td>
<td>...</td>
<td>1, 2, 4</td>
</tr>
<tr>
<td>HIP 29964</td>
<td>AO Men</td>
<td>K5.4</td>
<td>38.6 ± 1.3</td>
<td>9.77</td>
<td>15</td>
<td>4250$^{+140}_{-130}$</td>
<td>0.27 ± 0.04</td>
<td>370 ± 40</td>
<td>2.3</td>
<td></td>
</tr>
<tr>
<td>BD −17 6128 A</td>
<td>HD 358623 A</td>
<td>K5.7</td>
<td>45.7 ± 1.6$^a$</td>
<td>10.6$^b$</td>
<td>14</td>
<td>4130$^{+130}_{-130}$</td>
<td>0.25 ± 0.025$^c$</td>
<td>410 ± 20</td>
<td>1.2, 2, 5</td>
<td></td>
</tr>
<tr>
<td>CD −64 1208</td>
<td></td>
<td>K7.2</td>
<td>28.6 ± 0.2</td>
<td>9.54</td>
<td>110</td>
<td>4020$^{+140}_{-140}$</td>
<td>0.21 ± 0.04</td>
<td>460 ± 40</td>
<td>1.2</td>
<td></td>
</tr>
<tr>
<td>HIP 23309</td>
<td>CD −57 1054</td>
<td>K7.8</td>
<td>26.8 ± 0.8</td>
<td>10.01</td>
<td>8</td>
<td>3890$^{+110}_{-100}$</td>
<td>0.15 ± 0.03</td>
<td>370 ± 50</td>
<td>1.2, 3, 4</td>
<td></td>
</tr>
<tr>
<td>AU Mic</td>
<td></td>
<td>M0.7</td>
<td>9.9 ± 0.1</td>
<td>8.81</td>
<td>9</td>
<td>3750$^{+120}_{-110}$</td>
<td>0.07 ± 0.01</td>
<td>120 ± 60</td>
<td>−0.1 ± 0.7</td>
<td>1, 2, 3</td>
</tr>
<tr>
<td>GJ 3305</td>
<td></td>
<td>M1.0</td>
<td>29.4 ± 0.3</td>
<td>10.59</td>
<td>5</td>
<td>3710$^{+120}_{-130}$</td>
<td>0.13 ± 0.02</td>
<td>110 ± 60</td>
<td>−0.4 ± 0.7</td>
<td>3</td>
</tr>
<tr>
<td>BD −17 6128 B</td>
<td>HD 358623 B</td>
<td>M2.9</td>
<td>45.7 ± 1.6$^d$</td>
<td>...</td>
<td>20$^d$</td>
<td>3390$^{+100}_{-100}$</td>
<td>0.05 ± 0.01</td>
<td>130 ± 60</td>
<td>−0.2 ± 0.7</td>
<td>5</td>
</tr>
<tr>
<td>HD 155555 C</td>
<td></td>
<td>M3.5</td>
<td>31.4 ± 0.5$^e$</td>
<td>12.71</td>
<td>6</td>
<td>3340$^{+100}_{-100}$</td>
<td>0.05 ± 0.01</td>
<td>&lt;30</td>
<td>−1.4, 2</td>
<td></td>
</tr>
<tr>
<td>GJ 799 N</td>
<td>AT Mic</td>
<td>M4.4</td>
<td>10.7 ± 0.4</td>
<td>11.02</td>
<td>10</td>
<td>3210$^{+100}_{-100}$</td>
<td>0.04 ± 0.01</td>
<td>&lt;10</td>
<td>−1.4, 2.3</td>
<td></td>
</tr>
<tr>
<td>GJ 799 S</td>
<td>AT Mic</td>
<td>M4.4</td>
<td>10.7 ± 0.4</td>
<td>11.02</td>
<td>16</td>
<td>3210$^{+100}_{-100}$</td>
<td>0.04 ± 0.02</td>
<td>&lt;10</td>
<td>−1.4, 2.3</td>
<td></td>
</tr>
</tbody>
</table>

Notes. Except where noted, distances are from the updated Hipparcos catalog (van Leeuwen 2007). $V$ magnitudes are from Zuckerman et al. (2001), and the other data are our own measurements. The given $v \sin i$ is the average of the literature values cited in the last column. (1) $v \sin i$ given in Torres et al. (2006), (2) $v \sin i$ given in Scholz et al. (2007), (3) Spectral type from Zuckerman et al. (2001), (5) Luminosity from Neuhäuser et al. (2002)

$^a$ Assuming that BD −17 6128 A and B have the same parallax as their companion HD 199143.

$^b$ Value for the combined system.

$^c$ We give the corrected uncertainty in the luminosity of BD −17 6128 A (R. Neuhäuser 2009, private communication)

$^d$ The $v \sin i$ of BD −17 6128 B is not given in the literature. The value given here was estimated from a rotationally broadened stellar template (Section 3).

$^e$ Assuming that HD 155555 C is at the same distance as HD 155555 A/B.

of growth given in Soderblom et al. (1993) and Palla et al. (2007) to convert the measured EWs to lithium abundances, following the method used by Sestito et al. (2008). We assumed log $g = 4.5$, which is appropriate for late-type stars descending onto the main sequence. We linearly interpolated the curves of growth in $A$(Li)−EW(Li$i$) space and used a linear extrapolation for abundances <0.0 dex. For the three stars with effective temperature 3600 K < $T_{\text{eff}}$ < 4000 K, we linearly interpolated between the two tables. For the stars for which we have only an upper limit to the EW, the calculated abundance is necessarily also an upper limit. The quoted uncertainties are the result of individually perturbing both the effective temperatures and EW(Li$i$s) within their uncertainties and adding the effects on the derived abundances in quadrature. The results are given in Table 1.

Because HD 155555 A/B is an SB, we were not able to determine its lithium abundance from the EW(Li$i$s). The height of the continuum is roughly twice what would be measured for each star individually; therefore, the lithium EWs of the individual stars would be larger than the EWs given in Table 1, which are measured from the combined spectrum. We use the abundances given in Randich et al. (1993) of 3.5 for HD 155555 A and 3.7 for HD 155555 B. Based on the discussion in that paper and in Pasquini et al. (1991), we assume an uncertainty of ±0.3 dex on those abundances.

3.4. Lithium Depletion Boundary

We did not detect the LDL of the BPMG. The latest-type stars in our sample, HD 155555 C, GJ 799 N, and GJ 799 S, showed no observable lithium lines with very sensitive upper limits (Figure 1). The LDL of the BPMG has been detected in two binary systems that are members of this group: HIP 112312 A/B and V343 Nor A/B (Song et al. 2002, 2003). The 6707.8 Å Li line was detected in the lower-mass stars in these binaries whose spectral types are M4.5 and M5, respectively. The stars in our sample with non-detections of the lithium line have slightly earlier spectral types, which is consistent with the detected LDL. Comparison of the data with evolutionary models (see below; Figure 3) implies an LDL age of >30 Myr. Our data support the conclusion of Song et al. (2002) that there is a discrepancy between the LDL age of this group and the proposed age from H–R diagram isochrones of 12+8 Myr (Zuckerman et al. 2001).

3.5. Comparison to Other Work

To date, two other groups have published lithium abundances and effective temperatures for some members of the BPMG (Torres et al. 2006; Mentuch et al. 2008). Given the variation in effective temperatures and lithium abundances measured by each group, our data are consistent with previous measurements. Overall, we see the same trend as other groups with lithium decreasing as a function of decreasing spectral type and only upper limits in lithium abundance measured for GJ 799 N and S.

3.6. Comments on Individual Objects

3.6.1. CD −64 1208

Examination of the 6707.8 Å Li$i$ line and the Ca$i$ line at 6717.7 Å suggests that CD −64°1208 is a double-lined SB (Figure 2). Cross-correlation of the spectrum with radial velocity (RV) standards of known velocity and similar spectral type gives velocities of ±28±5 km s$^{-1}$ for the secondary and 43±5 km s$^{-1}$ for the primary. Based on the height of the cross-correlation peaks, the secondary has roughly 70%–80% the luminosity of the primary. We note that Torres et al. (2006) also flag this system as “SB?,” i.e., a possible spectroscopic binary. Our derived spectral type of K7 agrees with that of Riaz et al. (2006) and is similar to the K5 spectral type found by Torres et al. (2006). We note that the $v \sin i$ of 110 km s$^{-1}$ measured by Torres et al. (2006) is likely an overestimate. The rotation is rapid, but some of the line broadening is due to the velocity separation of the two
stars. If the luminosities and effective temperatures of the two components are similar, as seems likely from the spectrum, the EWs of the individual lithium line and the lithium abundances of the individual components should be roughly the same as those given in Table 1.

3.6.2. AU Mic

If AU Mic is the same age as the other members of the BPMG, it appears to be underluminous in the H–R diagram (Figure 3). We observe \( \log(L/L_\odot) = -1.15 \), but from the figure we might expect a value closer to \( \log(L/L_\odot) = -0.95 \) for it to lie on the same H–R diagram isochrone as the other BPMG members. AU Mic is known to have an edge-on debris disk (cf. Kalas et al. 2004), which could lead to some extinction. Given the difference between our observed and expected values for the luminosity, if AU Mic is coeval on the H–R diagram, then we estimate an extinction due to the debris disk of \( A_V \sim 0.5 \) mag. Since this extinction is not well characterized, and it is dependent on an assumption of coevality, we do not take it into account in our analysis. This proposed extinction should not affect our measurement of the effective temperature since that is determined from the spectral type which is measured over a small range in wavelength. However, it does affect any age determined from the H–R diagram position. If the luminosity we observe is less than the intrinsic luminosity of AU Mic, we would overestimate the age of the star (see Figure 3).

4. COMPARISON TO EVOLUTIONARY MODELS

In Figure 3, we plot our data against three sets of evolutionary models: D’Antona & Mazzitelli (1997, 1998), Baraffe et al. (1998), and Siess et al. (2000). The Siess et al. (2000) models include the effects of convective overshooting. Known multiple systems are plotted in color. We also plot HIP 112312 A/B and V343 Nor A/B, two systems that cross the LDB. We calculated luminosities and effective temperatures for these stars in the same manner as for our sample, using published spectral types and \( V \) magnitudes (Zuckerman et al. 2001; Song et al. 2003; Torres et al. 2006). We used the lithium abundances from Torres et al. (2006). Neither Song et al. (2002) nor Torres et al. (2006) quote a lithium abundance for HIP 112312 A. Since the spectral type and the upper limit in the EW of this star given in Song et al. (2002) are similar to those of GJ 799 N, we assign HIP 112312 A an upper limit of \( A(Li) = -1.4 \) dex.

To explore the differences between the models and our data, we independently determine ages for each star in two ways: from its position on the H–R diagram and from its position relative to the lithium depletion isochrones. We estimate the uncertainty of the H–R diagram age by perturbing the effective temperatures and luminosities within their uncertainties and adding the resulting age errors in quadrature. Because the measured lithium abundance is correlated with effective temperature, we estimate the uncertainties in the lithium depletion age by simultaneously varying the effective temperatures and lithium abundances within their uncertainties and taking the extreme values for the age as a measure of its uncertainty. These data are given in Table 2 and shown in Figure 4. This figure shows that the ages calculated from the H–R diagram are roughly consistent with a single age, but the ages measured from the lithium depletion have systematic trends. The trend of increasing age with spectral type found using the D’Antona & Mazzitelli (1997, 1998) and Siess et al. (2000) lithium depletion isochrones simply reflects
the fact that the data cross the isochrones instead of lying parallel to them, as seen in Figure 3. Although the Baraffe et al. (1998) models are the closest to predicting a coeval group using lithium depletion, they still fail to accurately represent the latest-type stars in the sample. Even though the predicted ages are similar, the pattern of lithium depletion seen in the Baraffe et al. (1998) models is the opposite of what is seen in our data. In particular, these lithium depletion isochrones indicate that the LDB occurs between 3400 K and 3200 K for the 30 Myr isochrone, which approximates our data. Thus, we expect that stars at the cooler end of the temperature range 3200–3400 K should have a higher lithium abundance than stars at the warmer end. In contrast, we observe that HD 155555 C, GJ 799 N, and GJ 799 S with $T_{\text{eff}} \sim 3300$ K, 3200 K, and 3200 K, respectively, have only upper limits in lithium abundance, whereas the warmer star BD $-17^h6128$ B ($T_{\text{eff}} \sim 3400$ K) has a larger lithium abundance.

From these two methods for measuring ages, we observe two systematic trends in the data. First, the ages of the mid-M dwarfs calculated from the H–R diagram isochrones are slightly younger than the ages of the warmer stars. This effect has also been noted by Stassun et al. (2004) in the ages of binary components and has been recognized more generally as a problem for PMS evolutionary models by Hillenbrand et al. (2008). The second trend is that the ages of the mid-M stars

<table>
<thead>
<tr>
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</thead>
<tbody>
<tr>
<td>HIP 29964</td>
<td>12$^{+6}<em>{-6}$ A (4$^{+1}</em>{-1}$)</td>
<td>31$^{+11}<em>{-11}$ A (40$^{+13}</em>{-13}$)</td>
<td>23$^{+9}<em>{-9}$ A (6$^{+2}</em>{-2}$)</td>
</tr>
<tr>
<td>BD $-17^h6128$ A</td>
<td>10$^{+5}<em>{-5}$ A (4$^{+1}</em>{-1}$)</td>
<td>26$^{+9}<em>{-9}$ A (35$^{+9}</em>{-9}$)</td>
<td>19$^{+8}<em>{-8}$ A (5$^{+3}</em>{-3}$)</td>
</tr>
<tr>
<td>HIP 23309</td>
<td>16$^{+8}<em>{-8}$ A (5$^{+1}</em>{-1}$)</td>
<td>29$^{+11}<em>{-11}$ A (28$^{+13}</em>{-13}$)</td>
<td>18$^{+9}<em>{-9}$ A (7$^{+1}</em>{-1}$)</td>
</tr>
<tr>
<td>CD $-64^d1208$</td>
<td>11$^{+6}<em>{-6}$ A (4$^{+1}</em>{-1}$)</td>
<td>26$^{+10}<em>{-10}$ A (29$^{+10}</em>{-10}$)</td>
<td>18$^{+9}<em>{-9}$ A (5$^{+1}</em>{-1}$)</td>
</tr>
<tr>
<td>AU Mic</td>
<td>32$^{+13}<em>{-13}$ A (9$^{+2}</em>{-2}$)</td>
<td>49$^{+15}<em>{-15}$ A (32$^{+12}</em>{-12}$)</td>
<td>28$^{+11}<em>{-11}$ A (11$^{+2}</em>{-2}$)</td>
</tr>
<tr>
<td>GJ 3305</td>
<td>9$^{+4}<em>{-4}$ A (9$^{+2}</em>{-2}$)</td>
<td>18$^{+6}<em>{-6}$ A (31$^{+10}</em>{-10}$)</td>
<td>10$^{+3}<em>{-3}$ A (12$^{+3}</em>{-3}$)</td>
</tr>
<tr>
<td>BD $-17^h6128$ B</td>
<td>11$^{+5}<em>{-5}$ A (17$^{+2}</em>{-2}$)</td>
<td>16$^{+7}<em>{-7}$ A (22$^{+6}</em>{-6}$)</td>
<td>13$^{+4}<em>{-4}$ A (24$^{+7}</em>{-7}$)</td>
</tr>
<tr>
<td>HD 155555 C</td>
<td>9$^{+6}_{-6}$ A ( &gt; 22)</td>
<td>13$^{+7}_{-7}$ A ( &gt; 32)</td>
<td>11$^{+5}_{-5}$ A ( &gt; 31)</td>
</tr>
<tr>
<td>GJ 799 N</td>
<td>7$^{+5}_{-5}$ A ( &gt; 31)</td>
<td>6$^{+5}_{-5}$ A ( &gt; 43)</td>
<td>8$^{+4}_{-4}$ A ( &gt; 36)</td>
</tr>
<tr>
<td>GJ 799 S</td>
<td>7$^{+5}_{-5}$ A ( &gt; 32)</td>
<td>6$^{+5}_{-5}$ A ( &gt; 44)</td>
<td>8$^{+4}_{-4}$ A ( &gt; 36)</td>
</tr>
</tbody>
</table>

Figure 4. Age as a function of spectral type, calculated for individual stars from different models. The left column shows ages calculated from the H–R diagram and the right column shows ages calculated from the lithium depletion isochrones. Since GJ 799 N and S have the same spectral type, the points for GJ 799 S have been displaced by +0.1 in spectral type for clarity.
The multiplicity frequency of stars at the youngest observable PMS ages is comparable to or greater than that among field stars (Duchêne et al. 2007, and references therein), indicating that binary formation is part of the star formation process. Thus, we can consider binary systems to be $N = 2$, coeval clusters to within $<1 \text{ Myr}$, although there are hints of smaller age differences (Stassun et al. 2008).

5. LARGER M-DWARF RADII TO THE RESCUE?

We have shown that the pattern of Li depletion seen in the BPMG stars is at odds with the predictions of PMS Li depletion models, in the sense that the models underpredict the amount of Li depletion in the M stars in our sample. The evolutionary models have a certain amount of intrinsic uncertainty due to assumptions about mixing length, equation of state, opacity, and atmosphere and boundary conditions, whose effects on lithium depletion are discussed in some detail in Burke et al. (2004).

Without a specific physical reason to change the assumptions, it is difficult to decide which parameter or combination of parameters should be changed in order to fit the lithium data. However, there is an increasingly large body of observations (see below) that evolutionary models tend to underpredict the radii of low-mass stars, particularly M stars, which may point the way toward the types of modifications to the models necessary to bring the H–R diagram ages and Li depletion ages into agreement.

Changes in stellar radii could affect the observed pattern of Li depletion in two ways: by actually changing Li depletion as a function of mass, and/or by altering the effective temperatures of stars, thus shifting them relative to the model isochrones. 

Regarding actual changes in Li depletion, the work of King et al. (2010) suggests a connection between Li depletion and stellar radii. They propose that a range in radii at a given stellar mass (perhaps due to a range in rotation rates and/or chromospheric activity) would lead to a range of interior temperature profiles and could explain the dispersion in lithium depletion in the Pleiades.

However, even if the models are essentially correct in terms of the interior conditions for a given mass, which largely govern the rate of Li depletion, there could still be a disagreement between the model predictions and the observations if the models do not correctly predict the surface conditions, e.g., the radius or effective temperature. We suggest here that current data show evidence of exactly this sort of disagreement.

There is mounting observational evidence that main-sequence evolutionary models predict radii for M dwarfs that are 10%–20% smaller than the observed radii. This effect is seen in stellar angular diameters measured interferometrically (Lane et al. 2001; Ségransan et al. 2003; Berger et al. 2006), in radii determined from eclipsing binaries (López-Morales 2007; Ribas et al. 2008; Fernandez et al. 2009), and in radii determined from detailed modeling of multi-band photometry to measure $L_{\text{bol}}$ and $T_{\text{eff}}$ (Mullan & MacDonald 2001; Casagrande et al. 2008).

There is still some question as to whether this discrepancy exists only for stars that are the most magnetically active (Demory et al. 2009), or if it applies to all M dwarfs (Casagrande et al. 2008). There does seem to be at least some differential effect from stellar activity (Mullan & MacDonald 2001; Morales et al. 2008), and models incorporating the effects of stellar activity are better able to reproduce the observed radii (Chabrier et al. 2007). For the present work, the question is moot as the stars in our sample (and indeed all late-type PMS stars) are very active, with large X-ray luminosities and active chromospheres.

The discrepancy between the true radii and the radii predicted by models very likely exists in the PMS phase as well, if it is due to effects of magnetic activity and/or missing opacity sources in the models. If this discrepancy is present, how would it manifest itself? Here, we argue that the effect would be similar to what we observe, and correcting for this would bring H–R diagram ages and Li depletion ages closer together. Corrected model isochrones for a given age would need to move to lower $T_{\text{eff}}$, increasing the age inferred from the H–R diagram and decreasing the age inferred from Li depletion for a given star.
First consider the H–R diagram. PMS stars change position in the H–R diagram as they age precisely because their radii are changing, as the stars descend (contract) toward the main sequence. If, at any given age, the true radii of stars are larger, then the model isochrones for that age need to move up and to the right. Thus, the H–R diagram age of a given star (of fixed, observed $T_{\text{eff}}$ and $L_B$/Mass) inferred from these corrected models would be older than before. Put another way, if radii are inflated by effects other than youth, it will take longer for a star to contract to a given radius, and thus to reach a given position on the H–R diagram.

Now, consider the $A$(Li)–$T_{\text{eff}}$ plane. The amount of Li depletion at a given age is primarily driven by a star’s mass, which sets its interior conditions. Inflating the radius of a star will cause that star to have a cooler surface, i.e., a lower $T_{\text{eff}}$. Thus, corrected model isochrones would need to shift to the right for a given $A$(Li) and age. Put another way, a star with an inflated radius will have a hotter interior than we might otherwise expect given its surface appearance, and thus will exhibit more Li depletion at a given age and $T_{\text{eff}}$ than would be predicted if the radius were not modified. This matches well with the enhanced Li depletion we observe for the cooler stars.

Calculating new PMS evolutionary models incorporating these effects is beyond the scope of this work, but we can explore these effects to a limited degree by introducing an ad hoc shift to the existing model isochrones. As a test of this idea, we shifted the model isochrones of Baraffe et al. (1998) and compared them to our observations of $A$(Li) and $T_{\text{eff}}$ for the BPMG stars. Based on the observations of M stars discussed above, we assumed that the larger radii are present only for stars cooler than $\sim$4200 K (Casagrande et al. 2008), and that the model radii should be inflated by 10%. As shown in Figure 5, these corrections to the models do indeed bring the inferred Li depletion and H–R diagram ages of the cooler stars into better agreement, though some discrepancies remain. Clearly, a more systematic treatment of this effect is needed, perhaps taking into account effects of different activity levels as well as different masses.

To close, we note that one implication of this result is that neither the LDB ages nor the H–R diagram ages of M stars are correct, but that the answer lies somewhere between the two.

6. CONCLUSIONS

We have measured lithium abundances and effective temperatures for ten members of the BPMG. We compare these abundances to the predictions of the evolutionary models of D’Antona & Mazzitelli (1997, 1998), Baraffe et al. (1998), and Siess et al. (2000). We find that while the H–R diagram isochrones of these models reproduce our data fairly well, the lithium depletion isochrones do not. In particular, the models predict less lithium depletion for the latest-type stars than we observe. Thus, the ages determined based on the lithium abundances are older than the ages determined from the H–R diagram.

The most striking result is the non-detection of lithium in the latest-type members (M3.5–M5) of this group. The lack of an observable lithium line in HD 155555 C, GI 799 N, and GI 799 S indicates that lithium depletion proceeds at a much faster rate than is predicted in the theoretical models given their spectral types. One potential clue lies in the discrepancy between the observed radii of M dwarfs and the radii predicted by models.

Regardless of the physical mechanism, it appears that mid-M stars deplete their lithium rapidly. The presence of the lithium line at 6707.8 Å in a star of this type remains a clear indicator of youth and suggests an age younger than that of the BPMG, while the lack of a detectable line is not a strong indicator that a star is old.

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